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Energy and energy flux in axisymmetric slow and fast waves

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ABSTRACT

Aims. We aim to calculate the kinetic, magnetic, thermal, and total energy densities and the flux of energy in axisymmetric sausage modes. The resulting equations should contain as few parameters as possible to facilitate applicability for different observations.

Methods. The background equilibrium is a one-dimensional cylindrical flux tube model with a piecewise constant radial density profile. This enables us to use linearised magnetohydrodynamic equations to calculate the energy densities and the flux of energy for axisymmetric sausage modes.

Results. The equations used to calculate the energy densities and the flux of energy in axisymmetric sausage modes depend on the radius of the flux tube, the equilibrium sound and Alfvén speeds, the density of the plasma, the period and phase speed of the wave, and the radial or longitudinal components of the Lagrangian displacement at the flux tube boundary. Approximate relations for limiting cases of propagating slow and fast sausage modes are also obtained. We also obtained the dispersive first-order correction term to the phase speed for both the fundamental slow body mode under coronal conditions and the slow surface mode under photospheric conditions.

Key words. Magnetohydrodynamics (MHD) - Methods: analytical - Sun: atmosphere - Sun: oscillations

1. Introduction

The solar atmosphere is a dynamic, magnetised plasma containing numerous distinct structures from active regions to coronal loops. These structures have been observed to support different oscillatory magnetohydrodynamic (MHD) modes (for recent reviews see Banerjee et al. 2007; De Moortel 2009; De Moortel & Nakariakov 2012). Magnetohydrodynamic waves are important for several reasons. First, they can be used to calculate the important background plasma parameters using MHD seismology of the solar atmosphere (e.g. Andries et al. 2005; Banerjee et al. 2007; Andries et al. 2009; De Moortel & Nakariakov 2012). Second, because MHD waves can carry energy over large distances, it has historically been thought that they can play a major role in the heating of the corona. Extensive discussions on coronal heating can be found in Walsh & Ireland (2003), Erdélyi (2004), Ofman (2005), and numerous others. To quantify the contribution of MHD waves to coronal heating we need both observations of MHD waves and methods to estimate the energy of those waves.

There have been many observations of MHD waves in different layers of the solar atmosphere with different instruments. In coronal loops MHD waves were observed with the Extreme ultraviolet Imaging Telescope (EIT) aboard the Solar and Heliospheric Observatory (SOHO) (e.g. Thompson et al. 1999), with the Coronal Multi-Channel Polarimeter (CoMP) (e.g. Tomczyk et al. 2007), with the extreme ultraviolet imaging spectrometer (EIS) aboard Hinode (e.g. Van Doorsselaere et al. 2008; Banerjee et al. 2009), and with the Solar Dynamics Observatory Atmospheric Imaging Assembly (SDO/AIA) (e.g. Aschwanden & Schrijver 2011; Morton et al. 2012b) to give but a few examples. In the chromosphere, MHD waves have been detected using the Solar Optical Telescope (SOT) aboard Hinode (De Pontieu et al. 2007), using the Swedish Solar Telescope (SST) (e.g. Jess et al. 2009), and using the Rapid Oscillations in the Solar Atmosphere (ROSA, Jess et al. 2010) instrument (e.g. Morton et al. 2011). In sunspots MHD waves have also been observed in smaller magnetic pores in the photosphere (e.g. Dorotovič et al. 2008, 2014; Yuan et al. 2014; Moreels et al. 2015). This non-exhaustive list of observations shows that there is a need for energy quantification methods in order to understand the energy supply of compressive MHD waves to the higher layers of the solar atmosphere. In these layers the wave energy is crucial to understanding the coronal heating problem.

There have been several papers that have estimated the energy in transverse kink oscillations. Morton et al. (2014) have used the ubiquity of propagating kink mode observations in the chromosphere to estimate the energy content of the observed kink waves. Thurgood et al. (2014) have detected transverse motions in solar plumes and used the displacement to estimate the energy in these transverse motions. Determining the energy of transverse kink motions has also received theoretical attention. Goossens et al. (2013a) used MHD theory to derive formulas that determine the energy of kink waves in the solar corona. Van Doorsselaere et al. (2014) have expanded on this idea and used MHD theory to write an expression for energy flux in terms of density filling factors.

In this paper we focus on axisymmetric slow and fast sausage modes, which have recently been discovered in numerous structures, but have not been adequately modelled to estimate their energy content. They have been detected in coronal loops (De Moortel et al. 2000; Wang et al. 2002, 2003; Krishna Prasad et al. 2012), in photospheric pores (Dorotovič et al. 2008, 2014;
Morton et al. 2011; Moreels et al. 2015), in network bright points (McAteer et al. 2003; Bharti et al. 2006; Martínez González et al. 2011), in flares (Nakariakov et al. 2003; Melnikov et al. 2005; Van Doorsselaere et al. 2011), and in the chromosphere (Morton et al. 2012a). Some authors have attempted to estimate the amount of energy present in these observations. Morton et al. (2012a) have observed both sausage and kink modes in the chromosphere and used MHD theory to estimate the energy in the waves. They found that the energy of compressive (sausage) waves was higher than the energy in transverse waves. We believe that the approach used in Morton et al. (2012a) can be improved in several ways. First, the integration of the wave energy should be performed over the entire cylinder since the eigenfunctions change with radial position. Second, the energy outside the flux tube should also be calculated since this can be of the same order of magnitude as the internal energy.

In this work we calculate the wave energy in sausage modes using a similar approach to that presented by Goossens et al. (2013a). The aim is to calculate the energy using only the background plasma equilibrium parameters and the phase speed of the wave. In this way the formulas can easily be applied to different observations to estimate the energy of sausage modes. This is a very important aspect since it means that energetics can be determined at height-localised regions of the solar atmosphere. In combination with the data analysis method described in Moreels et al. (2015) we do not need multiwavelength imaging approaches to determine the phase speed of the wave. Therefore, we are able to determine the energy in sausage oscillations for single filter measurements.

The paper is organised as follows: Sect. 2 gives the mathematical framework; Sect. 3 describes some limiting cases for the energy calculation which are of physical interest, i.e. slow sausage modes in coronal loops (Sect. 3.1 and Appendix A) and fast sausage modes in coronal loops (Sect. 3.2); Sect. 4 lists the conclusions of this paper.

2. Mathematical framework

The mathematical framework in which the energy for sausage waves is calculated is split into three sections. Section 2.1 derives general energy equations which can be used for different waves modes in different atmospheres. In Sect. 2.2 we calculate the energy inside the flux tube for surface modes. Section 2.3 lists the energy equations for both the energy inside the flux tube for body modes and the energy outside the flux tube.

2.1. General energy equations

Our equilibrium model is a straight cylinder with a constant radius \( R \) where the plasma is uniform both inside and outside the cylinder with a possible jump at the boundary (see Edwin & Roberts 1983). The magnetic field is directed along the axis of the flux tube and is given by \( B_{0,i} \) inside the flux tube and \( B_{0,e} \) outside the flux tube. The plasma pressure and density are \( \rho_{0,i} \) and \( \rho_{0,e} \) inside the flux tube and \( \rho_{0,i} \) and \( \rho_{0,e} \) outside the flux tube. We assume that the plasma has no background flow, i.e. the equilibrium velocity is \( v_0 = 0 \) both inside and outside the flux tube (see Fig. 1).

Before starting the energy calculations it is instructive to first consider the limitations of this equilibrium model. The straight uniform flux tube model neglects several physical effects, e.g. density stratification, flux tube expansion, and flux tube curvature. These effects can be important depending on the solar structure under consideration. The density stratification combined with the flux tube expansion is very important in the photosphere/chromosphere region. However, while the uniform flux tube model neglects these effects it is still valid when used as a first-order approximation. Andries & Cally (2011) have shown that for a slowly expanding flux tube the perpendicular eigenfunctions of MHD waves remain very similar and are still in terms of Bessel functions. In the corona the density stratification and flux tube expansion are minimal, but the curvature of the flux tubes can be a significant effect. Although Van Doorsselaere et al. (2004b) have shown that the curvature has no effect on eigenfrequencies of kink modes for a particular density profile, Verwichte et al. (2006) have shown that curvature could result in leaky wave modes, especially for thick flux tubes. While the leaky nature of the MHD waves will drastically change the energy outside the flux tube it will have a minimal impact on the energy inside the flux tube assuming that the internal remaining energy only decreases slowly over one period of the wave. Therefore, the uniform flux tube model results for the energy inside the flux tube are still very useful for estimating the energy content inside the flux tube of slowly leaking MHD waves.

Zaitsev & Stepanov (1975), Edwin & Roberts (1983), Sakurai et al. (1991), and many others have solved the ideal MHD equations for this equilibrium configuration. The solution method uses the following steps. First, the ideal MHD equations are linearised around the equilibrium state, e.g. \( \tilde{p}(r,\varphi,z,t) = \rho_0 + \tilde{\rho}(r,\varphi,z,t) \), where \( \rho_0 \) is the background and \( \tilde{\rho}(r,\varphi,z,t) \) is the small perturbation. From now on we drop the \( \varphi \) dependence since we are studying axisymmetric wave modes. Second, the resulting equations are solved analytically by performing a Fourier analysis in the ignorable coordinates, e.g. the density perturbation inside the flux tube can be written as \( \tilde{\rho}(r,z,t) = \rho'(r) \exp[i(kz - \omega t)] \). In this way all perturbed quantities can be related with the radial component of the Lagrangian displacement \( \xi_r \) and the total Eulerian pressure perturbation \( P \). Before
listing the most important equations we would like to make the
distinction between Eulerian and Lagrangian variables. This
distinction can be easily explained from an observational point
of view. We assume that we have a series of intensity images
of the sun. When looking at a fixed set of pixels for all images we
are using the Eulerian intensity. On the other hand, a magnetic
structure can be used to define a set of pixels for each image that
describe the structure, i.e. the set of pixels can change at each
image. Defining the intensity based on this set of pixels means
using the Lagrangian intensity. More information on the differ-
ence between Eulerian and Lagrangian variables can be found in
Moreels et al. (2013).

Summarizing the most important equations, we have

\[ \frac{D}{dr} \left( r \xi \right) = -C P', \]

\[ \frac{d}{dr} P' = \rho_0 (\omega^2 - k^2 c_s^2) \xi, \]

where \( D \) and \( C \) are given by

\[ D = \rho_0 (c_s^2 + c_A^2)(\omega^2 - k^2 c_s^2), \]

\[ C = \omega^4 - k^2 (c_s^2 + c_A^2)(\omega^2 - k^2 c_s^2). \]

In the above equations we have used the Alfvén speed \( c_A \),
the sound speed \( c_s \), and the tube speed \( c_T = \left( \frac{c_s^2 + c_A^2}{c_s^2} \right)^{1/2} \).

The above equations are valid both inside and outside the flux tube.
Equations 1 and 2 can be combined to form an ordinary dif-
fential equation for the total Eulerian pressure perturbation
\( P' \), which is solvable in terms of Bessel functions. Other use-
fule equations are

\[ \xi = \frac{\text{ik} \omega^2}{\rho_0 (c_s^2 + c_A^2)} (\omega^2 - k^2 c_s^2) P', \]

\[ \xi = 0, \]

\[ B' = \text{ik} B_0 \xi - B_0 (\nabla \cdot \xi) 1, \]

\[ p' = -\rho_0 c_s^2 \nabla \cdot \xi, \]

\[ \nabla \cdot \xi = \frac{-\omega^2}{\rho_0 (c_s^2 + c_A^2)} (\omega^2 - k^2 c_s^2) P', \]

\[ E' = \text{i} \omega B_0 [\xi, \nabla]. \]

Equations 3 provide algebraic expressions for the perturbed quan-
tities \( \xi, \xi, B', p', \nabla \cdot \xi, \) and \( E' \) in terms of \( \xi \) and \( P' \).

Using the continuity of \( \xi \) and \( P' \) at the flux tube boundary, the
dispersion relations can be readily derived (see Edwin & Roberts
1983). As an example, Fig. 2 shows the dispersion diagram un-
der photospheric conditions. For more information on the dif-
ferent wave modes and their eigenfunctions, see Moreels & Van
Doorselaere (2013).

To calculate the energy in sausage modes we follow the
method by Goossens et al. (2013a) (see also the correction for
the typographic error in Goossens et al. 2013b). This method is
based on the method described in Section 9.3 of the book by
Walker (2005). The analysis by Goossens et al. (2013a) is valid
for MHD waves in a pressureless plasma. In the present paper
plasma pressure is taken into account. This introduces additional
terms related to the thermal energy. To calculate the energy we
use quantities that are averaged over a complete cycle of the
wave. The averaged kinetic energy \( \langle KE \rangle \), the averaged magne-
tic energy \( \langle ME \rangle \), the averaged thermal energy \( \langle IE \rangle \), the averaged
total energy \( \langle TE \rangle \), the averaged Poynting flux \( \langle S \rangle \), the averaged

\[ \langle KE \rangle = \frac{\rho_0}{4} (v \cdot v^*), \]

\[ \langle ME \rangle = \frac{1}{4 \mu_0} (B' \cdot B'^*), \]

\[ \langle IE \rangle = \frac{1}{4 \mu_0 c_s^2} (p' p^*), \]

\[ \langle TE \rangle = \langle KE \rangle + \langle ME \rangle + \langle IE \rangle, \]

\[ \langle S \rangle = \frac{1}{2} \text{Re} \left( \frac{E' \times B'^*}{\mu_0} \right), \]

\[ \langle T \rangle = \frac{1}{2} \text{Re} \left( p' p^* \right), \]

\[ \langle F \rangle = \langle S \rangle + \langle T \rangle, \]

where \( \mu_0 \) is the magnetic permeability and an asterix denotes the
complex conjugate of a quantity. We note that the expressions
describe energy density since the dimension of energy is J/m³.

The energy flux is expressed in W/m². When considering the energy
in sausage modes it can seem strange that the energy equations are
quadratic in nature for linear wave modes. However, the method described in Walker (2005) uses
only first-order reduced MHD equations. The energy equations
also contain no linear terms because these integrate to zero over
a period/wavelength. On the other hand, it is true that when the
observed waves are non-linear in nature, the linear energy
equations might no longer be valid. Therefore, we have calculated
the largest radial component of the Lagrangian displacement in
order to still be a linear MHD wave. It is straightforward to ob-
tain a link between \( \xi / R \) and \( \rho' / \rho_0 \) and this shows that \( \xi / R \)
can be of the order of a few percent. These small displacements are
observable (e.g, Moreels et al. 2015; Grant et al. 2015), but not
all MHD wave mode observations are linear (e.g. Morton et al. 2012a). When applying the linear energy equations to non-linear systems there can be an overestimate of the energy and energy flux owing to non-linear saturation of the system. Using Eqs. 3 we can rewrite Eqs. 4 to find

$$\langle KE \rangle = \frac{\rho_0 \omega^2}{4} (\xi \cdot \xi^*),$$

$$\langle ME \rangle = \frac{\rho_0 \omega^2}{4} \left( \frac{\omega^2_{\perp} \xi}{\omega_{\perp}^2} + \frac{\omega^2_{\parallel} \xi}{\omega_{\parallel}^2} \right) \left[ (\nabla \cdot \xi)(\nabla \cdot \xi^*) \right]$$

$$= \frac{\rho_0 \omega^2}{4} \left( \frac{\delta \xi}{\delta r} \frac{\delta \xi^*}{\delta r} \right),$$

$$\langle IE \rangle = \frac{\rho_0 \omega^2}{4} \left[ (\nabla \cdot \xi)(\nabla \cdot \xi^*) \right],$$

$$\langle TE \rangle = \langle KE \rangle + \langle ME \rangle + \langle IE \rangle,$$

$$\langle S \rangle = -\frac{1}{2} \text{Re} \left\{ \left( i \omega_1 \omega \xi_r^2 (\nabla \cdot \xi^*) \right) \xi^* + \left( \rho_0 \omega^2 \xi_r^2 \right) \xi^* \right\},$$

$$\langle T \rangle = -\frac{1}{2} \text{Re} \left\{ i \omega \rho_0 \xi_r^2 (\nabla \cdot \xi^*) \right\},$$

$$\langle E \rangle = \langle S \rangle + \langle T \rangle.$$  

Equations 5 allow us to calculate the energy for any wave mode in a cylindrical plasma structure.

2.2. Energy inside the flux tube for surface modes

We now have three cases (see Fig. 2) in which to calculate the energy and energy flux. First, the energy inside the flux tube for a surface mode. Second, the energy inside the flux tube for a body mode. Third, the energy outside the flux tube. All three calculations are very similar and we only give the details of the calculation for the energy inside the flux tube for a surface mode. The Eulerian total pressure perturbation for a sausage surface mode is given by

$$P'(r, z, t) = A \omega I_0(\kappa r) \exp(i(kz - \omega t)),\quad (6)$$

where $A$ is a constant amplitude with dimension $J/m^3$, $I_0(\cdot)$ indicates the modified Bessel function of the first kind of order zero, and $\kappa_r$ is given by

$$\kappa_r^2 = k^2 \left( \frac{\omega^2_{\perp}}{\omega_{\perp}^2} - \frac{\omega^2_{\parallel}}{\omega_{\parallel}^2} \right) \left( \frac{\omega^2_{\perp}}{\omega_{\perp}^2} + \frac{\omega^2_{\parallel}}{\omega_{\parallel}^2} \right),\quad (7)$$

where $k^2$ is positive. We have introduced the internal sound frequency $\omega_{\parallel,i}$, the internal Alfvén frequency $\omega_{\perp,i}$, and the internal tube frequency $\omega_{T,i}$, which are given by $\omega_{\parallel,i} = k c_{i,i}$. We also need expressions for the radial and longitudinal components of the Lagrangian displacement. Using Eqs. 2 and 3 results in

$$\xi_r(r, z, t) = \Xi_r I_1(\kappa_r r) \exp(i(kz - \omega t)),\quad (8)$$

$$\xi_z(r, z, t) = \Xi_z I_0(\kappa_r r) \exp(i(kz - \omega t)),\quad (9)$$

where $\Xi_r$ and $\Xi_z$ are given by

$$\Xi_r = A \frac{\kappa}{\rho_0(\omega^2_{\perp} - \omega_{\perp}^2)},\quad (10)$$

$$\Xi_z = A \frac{ik \omega_{\perp,i}}{\rho_0(\omega^2_{\parallel,i} + \omega_{\parallel,i}^2)(\omega^2_{\perp} - \omega_{\perp}^2)}.\quad (11)$$

The constants $\Xi_r$ and $\Xi_z$ are related to the radial and the longitudinal components of the Lagrangian displacement at the flux tube boundary and have dimension m. Also clear is that $\Xi_r$ and $\Xi_z$ are related to each other by the amplitude $A$; therefore, quantifying the radial component of the Lagrangian displacement results in a measure for the longitudinal component of the Lagrangian displacement and the other way around as well. Later on we use the notation $\Xi'_r$ and $\Xi'_z$, which are the radial and the longitudinal components of the Lagrangian displacement at the flux tube boundary, respectively, and are given by

$$\Xi'_r = \max(\xi_r(R, z, t)) = \Xi_r I_1(\kappa R),\quad (12)$$

$$\Xi'_z = \max(\xi_z(R, z, t)) = \Xi_z I_0(\kappa R).\quad (13)$$

Calculating the magnetic energy also requires the divergence of the Lagrangian displacement, which is given in Eqs. 3 and results in

$$\nabla \cdot \xi = A \frac{i \omega^2_{\perp,i} - \omega^2_{\parallel,i}}{\rho_0(\omega^2_{\perp,i} + \omega_{\perp}^2)} I_0(\kappa r) \exp(i(kz - \omega t)).\quad (14)$$

Using Eqs. 6 - 14 we can now calculate the averaged energy and the averaged energy flux inside the flux tube (Eqs. 5), and we find

$$\langle KE \rangle = \frac{\rho_0 \omega^2}{4} \left( \left\langle |\xi|^2 \right\rangle + \left\langle |\xi|^2 \right\rangle \right),$$

$$\langle ME \rangle = \frac{\omega^2_{\perp,i}}{\omega^2_{\perp}} C_i I_1(\kappa_r r)^2 + C_z I_0(\kappa_r r)^2,$$

$$\langle IE \rangle = \frac{\omega^2_{\perp,i}}{\omega^2_{\perp}} C_i I_0(\kappa_r r)^2,$$

$$\langle S \rangle = 2 C_z \frac{\omega_{\perp,i}}{\omega} I_0(\kappa r)^2 1_z,$$

$$\langle T \rangle = 2 C_i \frac{\omega_{\perp,i}}{\omega} I_0(\kappa r)^2 1_z,$$

where $C_i$ and $C_z$ are constants and are given by

$$C_i = \frac{\rho_0 \omega^2}{4} \left\langle |\Xi|^2 \right\rangle,$$

$$C_z = \frac{\rho_0 \omega^2}{4} \left\langle |\Xi|^2 \right\rangle.$$
per unit length along the magnetic field. To calculate the total energy we should also integrate over the $z$-direction from a height zero to a height $L$, where $L$ is the length of the flux tube (see Van Doorsselaere et al. 2014). Since the wave mode properties do not change with height this results in a simple multiplication with a factor $L$. Closed analytic expressions for the energy and flux of energy can be found when using the following integrals:

$$\int_0^L uJ_0(u)^2 du = \frac{R^2}{2} (J_0(R)^2 + J_1(R)^2),$$

$$\int_0^L uJ_1(u)^2 du = \frac{R^2}{2} (J_1(R)^2 - J_0(J_1)J_2(R)).$$

These integrals can be checked with Eqs. 9.1.27 and 9.1.28 in Abramowitz & Stegun (1972). The result for the averaged energy and averaged energy flux for body modes is

$$\langle \mathbf{KE} \rangle = C_z(n,R)^2 \left( J_1(n,R)^2 - J_0(n,R)J_2(n,R) \right) + C_z(k,R)^2 \left( J_1(n,R)^2 + J_1(n,R)^2 \right),$$

$$\langle \mathbf{ME} \rangle = \frac{\omega^2_A}{\omega^2} \langle C_z(n,R)^2 \left( J_1(n,R)^2 - J_0(n,R)J_2(n,R) \right) \rangle + \omega^2_A \left( \frac{\omega_i - \omega}{\omega^2 - \omega^2_{\omega_i}} \right) \langle C_z(k,R)^2 \left( J_1(n,R)^2 + J_1(n,R)^2 \right) \rangle,$$

$$\langle \mathbf{IE} \rangle = \omega^2 \langle C_z(k,R)^2 \left( J_1(n,R)^2 + J_1(n,R)^2 \right) \rangle,$$

$$\langle \mathbf{TE} \rangle = \langle \mathbf{KE} \rangle + \langle \mathbf{ME} \rangle + \langle \mathbf{IE} \rangle,$$

$$\langle \mathbf{S} \rangle = 2\frac{\omega^2_A}{\omega^2} \langle C_z(n,R)^2 \left( J_1(n,R)^2 - J_0(n,R)J_2(n,R) \right) \rangle \mathbf{1}_z,$$

$$\langle \mathbf{T} \rangle = 2\frac{\omega^2}{k} \langle C_z(k,R)^2 \left( J_1(n,R)^2 + J_1(n,R)^2 \right) \rangle \mathbf{1}_z,$$

$$\langle \mathbf{F} \rangle = \langle \mathbf{S} \rangle + \langle \mathbf{T} \rangle, \quad (17)$$

where $C_z$ and $C_k$ are given by Eqs. 16.

In the same manner we can calculate the energy outside the flux tube. There are several changes in the derivation. First, the integration of the Bessel function changes, since externally the wave is written by $P(r,z,t) = A K_0(\kappa r) \exp(ikz - \omega t)$, where $A$ is a constant amplitude, $K_0(\cdot)$ indicates the modified Bessel function of the second kind of order zero, and $\kappa^2$ is the same as $k^2$ but with equilibrium parameters outside the flux tube. Second, the integrated quantities are now defined as

$$\langle \mathbf{KE} \rangle = 2\pi \int_\mathcal{R} \langle \mathbf{KE} \rangle rdr, \quad (18)$$

i.e. integrating over the entire domain outside the flux tube. To obtain closed analytic expressions for the energy and flux of energy we use Eqs. 9.6.26 and 9.6.28 in Abramowitz & Stegun (1972), i.e. we use

$$\int_\mathcal{R} uK_0(u)^2 du = \left[ \frac{1}{2} \left( \omega^2 K_0(u)^2 - \omega^2 K_1(u)^2 \right) \right]_R^\infty,$$

$$\int_\mathcal{R} uK_1(u)^2 du = \left[ \frac{1}{2} \left( \omega^2 K_1(u)^2 - \omega^2 K_0(u)K_2(u) \right) \right]_R^\infty.$$
the averaged energy flux for the energy outside the flux tube is
\[
\langle KE \rangle = C_1 (k_e R)^2 \left( K_0 (k_e R) K_2 (k_e R) - K_1 (k_e R)^2 \right) + C_2 (k R)^2 \left( K_1 (k R)^2 - K_0 (k R)^2 \right),
\]
\[
\langle ME \rangle = \frac{\omega^2_A}{\omega^2} C_1 (k_e R)^2 \left( K_0 (k_e R) K_2 (k_e R) - K_1 (k_e R)^2 \right),
\]
\[
\langle IE \rangle = \frac{\omega^2}{\omega^2} C_1 (k R)^2 \left( K_1 (k R)^2 - K_0 (k R)^2 \right),
\]
\[
\langle IE \rangle = \langle KE \rangle + \langle ME \rangle + \langle IE \rangle
\]
where \( C_1 \) and \( C_2 \) are given by Eqs. 16 but with all equilibrium parameters taken outside the flux tube.

Equations 15, 16, 17, and 19 allow us to calculate the energy in sausage mode in a general cylindrical plasma structure, i.e. we have not restricted ourselves to coronal flux tubes. As input we need the radius of the flux tube, the sound and Alfvén speeds, the plasma density, the period and phase speed of the wave, and the radial or longitudinal component of the Lagrangian displacement at the flux tube boundary. We stress that only the radial or the longitudinal component of the Lagrangian displacement is needed since both are linked through the amplitude \( A \) as mentioned earlier. All of these are available with modern ground- or space-based observations combined with empirical solar atmospheric models (e.g. recent work by Morton et al. 2011; Dorotović et al. 2014; Moreels et al. 2015). Our equations are generally valid and can be used to calculate the energy and energy flux in the observed axisymmetric waves.

The group speed \( v_g \) is defined as
\[
v_g = \frac{\langle F \rangle}{\langle IE \rangle},
\]
where \( \langle F \rangle \) is the sum of both the internal and external flux of energy and \( \langle IE \rangle \) also includes both internal and external energy. In this way the group speed expresses the propagation speed of energy.

### 3. Limiting cases

In this section we study some cases of physical interest, i.e. sausage modes in thin flux tubes. Figure 3 shows the dispersion diagram under coronal conditions. We have two types of wave modes, i.e. fast and slow sausage body modes. We only plotted one slow body mode and five fast body modes, while there are actually many radial overtones. We begin by studying slow sausage modes (Sect. 3.1) in thin (coronal) flux tubes and afterwards we look at fast sausage modes (Sect. 3.2) in coronal flux tubes at the cut-off wavenumber. The extra assumptions that are made in the section are usually valid in the solar corona, but care needs to be taken with the thin tube limit. Aschwanden et al. (2004) explained that coronal loops may oscillate with higher harmonics, which would result in the thin tube limit not being applicable, and that the use of the full energy equations from Sect. 2.2 should be applied. We stress that this section of limiting cases uses extra assumptions that are not necessarily valid in different layers of the solar atmosphere. Therefore, care needs to be taken when utilizing the results from limiting cases and in some cases it may be that the full set of energy equations 15 provide a more accurate result.

#### 3.1. Slow waves

In this section we look at the limiting case of slow waves in thin flux tubes. In the literature there has been extensive study on slow magnetoacoustic waves (see De Myrort (2009) for an overview). From these studies we know that the energy of the wave should propagate along the magnetic field at the local tube speed. We begin by giving the dispersion relation for slow sausage body modes (see Edwin & Roberts 1983),
\[
\rho_0 \left( \frac{\omega^2\lambda}{\omega^2} \right) K_0 (k_e R) J_0 (n R) = 1,
\]
where the dash denotes the derivative of a Bessel function (e.g. \( J_0'(n R) \)) and \( J_0(x) \) evaluated at \( x = n R \) and where \( n^2 = -k^2 \) is given by Eq. 7. We also study slow sausage surface modes, which do not occur under coronal conditions but do occur under photospheric conditions (see Fig. 2). The dispersion relation is slightly different, i.e.
\[
\rho_0 \left( \frac{\omega^2\lambda}{\omega^2} \right) K_0 (k_e R) J_0 (k R) = 1,
\]
but the calculations will not be very different. We study slow waves in thin flux tubes, meaning that \( k R \) goes to zero. When looking at the phase diagram (Fig. 2 or Fig. 3) we notice that the phase speed of slow waves approaches the internal tube speed as \( k R \) becomes smaller. After studying the dispersion relation (Eq.
21) we found that as $kR \to 0$ the frequency of body modes is given by $\omega = \omega_{1,i} + (kR)^2 \ln(kR) \omega_{1,j}$. For surface modes the $+$ sign is replaced with a $-$ sign. Expanding the Bessel functions that appear in the dispersion relation (Eq. 21) results in a non-zero $\omega_1$ given by

$$\omega_1 = \frac{1}{4} \frac{\rho_0 \varepsilon}{\rho_0} \left( \frac{\omega_{s,e}^2 - \omega_{s,i}^2}{\omega_{s,e}^2 + \omega_{s,i}^2} \right)^2 \frac{\omega_{s,e}^2 - \omega_{s,i}^2}{\omega_{s,e}^2 + \omega_{s,i}^2} \omega_{Tr,i}. \quad (23)$$

We want to make several remarks on this dispersive correction term $\omega_1(kR)^2 \ln(kR)$. First, the expression for $\omega_1$ is not needed in the long wavelength limit (i.e. $kR = 0$). Second, the expression for $\omega_1$ has the correct dimension, i.e. it is a frequency. Third, under coronal conditions the sign is negative as we would expect since slow body waves have higher phase speed than the internal tube speed. Fourth, this formula is only valid for the fundamental body mode, not for radial overtones because of the expansions of the Bessel functions we have used, i.e. we assumed that $nR \ll 1$, while for radial overtones this is not the case. Fifth, the formula for $\omega_1$ is not valid under photospheric conditions. Figure 3 in Moreels & Van Doorsselaere (2013) shows that the fundamental slow sausage body mode under photospheric conditions has one node inside the flux tube in the eigenfunctions, meaning that $nR$ is not small and our expansions of the Bessel functions are not valid. Because of these restrictions, this formula is only valid for the fundamental body mode under coronal conditions.

In expanding the Bessel functions we used

$$nR = \frac{1}{\ln(kR)} \frac{\left( \omega_{s,e}^2 - \omega_{s,i}^2 \right) \left( \omega_{s,i}^2 - \omega_{s,e}^2 \right)}{\left( \omega_{s,e}^2 + \omega_{s,i}^2 \right) (2\omega_{s,e} \omega_{s,i})} \to 0,$$
$$\kappa_e = kR \frac{\left( \omega_{s,e}^2 - \omega_{s,i}^2 \right) \left( \omega_{s,i}^2 - \omega_{s,e}^2 \right)}{\left( \omega_{s,e}^2 + \omega_{s,i}^2 \right) (2\omega_{s,e} \omega_{s,i})} \to 0,$$

where we retained terms up to $(kR)^2 \ln(kR)$. The constants $C_i$ and $C_e$ are written in several forms; one of which uses unknown amplitude $A$. We can find the amplitude $A$ by looking at the Lagrangian displacement and the pressure perturbation (Eqs. 6, 8, and 9). We know that all these must remain finite, also at the flux tube boundary. We have defined $\Xi'$ to indicate the maximum longitudinal component of the Lagrangian displacement at the flux tube boundary. This is related to the variable $\Xi_e'$ by $\Xi_e' = \Xi_e J_0(nR)$, meaning that for thin flux tubes we have $\Xi_e' = \Xi_e$. Assuming that $\Xi_e'$ is finite, combined with the continuity of the pressure perturbation allows us to find an expression for both the internal amplitude $A_i$ and the external amplitude $A_e$. These are given by

$$A_i = \frac{(2i\rho_0 \omega_{Tr,i} \omega_1)}{\rho_0} (1 + \omega_{s,i}^2/\omega_{s,e}^2) RE_e',$$
$$A_e = \frac{(2i\rho_0 \omega_{Tr,i} \omega_1)}{\rho_0} (1 + \omega_{s,i}^2/\omega_{s,e}^2) RE_e'.$$

We note that these are just Eqs. 10 and 11 applied to body modes in thin flux tubes. When using these amplitudes we find that both the pressure perturbation and the radial component of the Lagrangian displacement inside the flux tube go to zero as $(kR) \ln(kR)$ goes to zero. The longitudinal component of the Lagrangian displacement inside the flux tube is approximately constant and is non-zero. The pressure perturbation, the radial component of the Lagrangian displacement, and the longitudinal component of the Lagrangian displacement outside the flux tube go to zero as $kR$ goes to zero.

We now calculate the energy for slow sausage modes in the long wavelength limit (i.e. $kR = 0$) and we find

$$\langle KE \rangle = \frac{\rho_0}{4} \omega_{s,i}^2 \pi R^2 \Xi_e'^2,$$
$$\langle ME \rangle = \frac{\rho_0}{4} \omega_{s,i}^2 \pi R^2 \Xi_e'^2,$$
$$\langle IE \rangle = \frac{\rho_0}{4} \omega_{s,i}^2 \pi R^2 \Xi_e'^2,$$
$$\langle TE \rangle = \frac{\rho_0}{2} \pi R^2 \Xi_e'^2,$$
$$\langle S \rangle = 0,$$
$$\langle T \rangle = \frac{\rho_0}{2} \pi R^2 \Xi_e'^2 c_{Tr,i} \mathbf{1}_z,$$
$$v_g = c_{Tr,i} \mathbf{1}_z. \quad (24)$$

As before, the energy is expressed in J/m and the energy flux is expressed in W. The energy is contained entirely inside the flux tube, i.e. the energy outside the flux tube is zero. This was to be expected since we already found that the perturbations outside the flux tube go to zero as $kR$ goes to zero. We also retrieve the expected equipartition of energy between the potential energy (i.e. the sum of both magnetic and thermal energy) and the kinetic energy. The group speed is exactly equal to the tube speed and is directed along the magnetic field. We remark that these formulas are applicable in different regions of the solar atmosphere, both in the corona and in the photosphere. Previously, we stressed that the dispersive correction $\omega_1(kR)^2 \ln(kR)$ was only valid for fundamental body modes under coronal conditions. However, as mentioned before, in deriving Eqs. 24 we did not need Eq. 23, which describes $\omega_1. The constant $nR$ cancelled out in the long wavelength limit, showing that Eqs. 24 are also valid under photospheric conditions. In Appendix A we list the complete expressions for slow wave modes in thin flux tubes, i.e. with a small but non-zero value of $kR. The expressions in Appendix A are valid for the fundamental slow body mode under coronal conditions and for the surface mode under photospheric conditions.

We averaged the flux of energy across a surface perpendicular to the magnetic field resulting in $\langle F_S \rangle$. The surface over which we averaged is a circular disk with radius $d_t$. The idea behind this averaging over an area is linked to the paper by Van Doorsselaere et al. (2014). In that paper MHD theory is used to write an expression for energy flux in multistrand coronal loops in terms of density filling factors. As explained in Van Doorsselaere et al. (2014) the radius of a flux tube strand $R$ can be linked with the influence radius of that flux tube strand $d_t$ using the density filling factor $f$, i.e.

$$f = \frac{R^2}{d_t^2}. \quad (25)$$

The averaging results in

$$\langle F_S \rangle = \frac{\rho_0}{2} \pi R^2 \Xi_e'^2 c_{Tr,i} \mathbf{1}_z,$$
$$v_g = c_{Tr,i} \mathbf{1}_z. \quad (26)$$

where the averaged flux is now dependent on the density filling factor $f$. We note that the dimension of $\langle F_S \rangle$ is W/m$^2$.

For slow waves in the footpoints of thin coronal flux tubes we can make the extra assumption that the plasma beta is low
(i.e. $\beta \ll 1$). We find that $\omega_{j,1}^2 \approx \omega_{j,1}^2 (1 - 0.5 \gamma \beta)$, where $\gamma$ is the ratio of specific heats. For the energy we find

$$\langle KE \rangle = \frac{D_{0j}}{4} \omega_{j,1}^2 R^2 \Xi_c^2,$$

$$\langle ME \rangle = \frac{D_{0j}}{4} \omega_{j,1}^2 R^2 \Xi_c^2 \frac{\gamma \beta}{2},$$

$$\langle TE \rangle = \frac{D_{0j}}{4} \omega_{j,1}^2 R^2 \Xi_c^2 \left(1 - \frac{\gamma \beta}{2}\right).$$

$$\langle S \rangle = 0,$$

$$\langle T \rangle = 2 \frac{D_{0j}}{4} \omega_{j,1}^2 R^2 \Xi_c^2 c_{s1}1_z,$$

$$\nu_b = c_{s1} \left(1 + \frac{\gamma \beta}{2}\right) 1_z.$$  \hspace{1cm} (27)

When the plasma beta is equal to zero we still have equipartition between kinetic and potential energy, but the potential energy is only thermal energy with no contribution of magnetic energy. We also find that energy is transported along the field lines at the internal tube speed, which is equal to the internal sound speed. Therefore, these waves are very like sound waves in a non-magnetised atmosphere, except these waves travel along the magnetic field.

### 3.2. Fast waves

In this section we study fast sausage modes under coronal conditions. The fast sausage modes can be seen in the coronal dispersion diagram (Fig. 3) and their dispersion curves are shown with full lines. In particular we study the fundamental fast sausage mode, but radial overtones are also discussed. We immediately put plasma beta equal to zero since this eliminates only the slow modes from the dispersion diagram, not the fast modes. We split the discussion into two parts, one dealing with the energy inside the flux tube and the other outside the flux tube.

#### 3.2.1. Energy inside the flux tube

The dispersion relation for body modes has already been given in Eq. 21. Following the same calculation as Vasheghani Farahani et al. (2014) we find that the solution of the dispersion relation at the cut-off wavenumber $k_c$ is given by

$$\omega = \omega_{A,e}, n_R = j_{0,1},$$

where $j_{0,1}$ is the first zero of the Bessel function $J_0(n_R)$. This is the solution for the fundamental sausage mode; for the $\ell$th radial overtone we need to replace $j_{0,1}$ with $j_{0,\ell}$, where $j_{0,\ell}$ is the $\ell$th zero of the Bessel function $J_0(n_R)$.

Unlike the study of slow sausage waves, here we do not need a better approximation of the phase speed around the cut-off wavenumber since no indeterminate forms occur in the energy calculations. We know that

$$n_R = j_{0,1},$$

$$k_c R = k_{e,R} \frac{\left(\omega_{A,e}^2 - \omega_{A,e}^2\right) \left(\omega_{A,e}^2 - \omega_{e,e}^2\right)}{\left(\omega_{A,e}^2 + \omega_{A,e}^2\right) \left(\omega_{e,e}^2 - \omega_{A,e}^2\right)} = 0.$$

We can now expand the Bessel functions that occur in the energy equations (Eqs. 17). As was the case for slow modes, here too we need an expression for the unknown amplitude $A$ that appears in the constants $C_1$ and $C_2$. Again, we determine the amplitude $A$ from Eqs. 10 and 11 applied to body modes. Since we are working in the cold plasma approximation we already know that $\xi_c = 0$ everywhere, therefore Eq. 11 cannot be used. As before, we define $\Xi'_c$ as the maximum value of the radial component of the Lagrangian displacement at the flux tube boundary, which is linked to $\Xi_c$ by $\Xi'_c = \Xi_c J_1(n_R)$. We know that $\xi_c$ and $P'$ must be finite everywhere and non-zero at some point inside the flux tube, otherwise there would be no wave energy inside the flux tube. This happened in the previous limiting case (see Sect. 3.1), where in the long wavelength limit the external wave perturbations were exactly zero and therefore we found no external energy. A possible solution is

$$A_i = - \frac{\rho_{0,i} \left(\omega_{A,e}^2 - \omega_{A,e}^2\right)}{j_{0,1} J_1(j_{0,1})} R \Xi'_c,$$

$$A_e = - \frac{\rho_{0,e} \left(\omega_{A,e}^2 - \omega_{A,e}^2\right)}{(k_c R) J_1(k_c R)} R \Xi'_c = 0.$$

After some straightforward calculations we find that the energy inside the flux tube for fast sausage modes at the cut-off wavenumber in the zero plasma beta approximation is given by

$$\langle KE \rangle = \frac{D_{0j}}{4} \omega_{A,e}^2 R^2 \Xi_c^2,$$

$$\langle ME \rangle = \frac{D_{0j}}{4} \omega_{A,e}^2 R^2 \Xi_c^2,$$

$$\langle TE \rangle = 0,$$

$$\langle S \rangle = 2 \frac{D_{0j}}{4} \omega_{A,e}^2 c_{s1} R^2 \Xi_c^2 1_z,$$

$$\langle T \rangle = 0.$$ \hspace{1cm} (28)

Because we are working in the cold plasma approximation we expect no thermal energy and this is confirmed here. We note that there is no equipartition of energy between magnetic and kinetic energy. This is not a problem since we have not included the external energy at the moment. We have not mentioned the group speed since we have not yet calculated the external energy.

#### 3.2.2. Energy outside the flux tube

We now calculate the external energy for fast sausage modes at the cut-off wavenumber. From Sect. 3.2.1 we already know the complete solution to the dispersion relation and the amplitude $A_e$. Therefore, we can just use Eqs. 19. We know that the kinetic energy is given by

$$\langle KE \rangle_e = C_e (k_c R)^3 \left(K_0(k_c R) - K_1(k_c R)\right)^2$$

$$+ C_e (k_c R)^2 \left(K_1(k_c R)^2 - K_0(k_c R)^2\right)$$

$$+ \frac{D_{0j}}{4} \omega_{A,e}^2 R^2 \Xi_c^2 \left(-\ln(k_c R) - 1\right).$$

In the calculation we have used the small argument expansions for the Bessel functions since $\xi_c R = 0$. The external kinetic energy resembles the internal kinetic energy, except that the density is taken outside and there is the added $(-\ln(k_c R) - 1)$ term. It is this last term that is problematic: we know that $\kappa_c R = 0$, and so this term becomes infinitely large.

To expand on this case of infinite energy, we first looked at trapped wave modes very close to the cut-off wavenumber this term becomes infinitely large. We can now expand the Bessel functions that occur in the energy equations (Eqs. 17). As was the case for slow modes, here too we need an expression for the unknown amplitude $A$ that appears in the constants $C_1$ and $C_2$. Again, we determine the amplitude $A$ from Eqs. 10 and 11 applied to body modes. Since we are working in the cold plasma approximation we already know that $\xi_c = 0$ everywhere, therefore Eq. 11 cannot be used. As before, we define $\Xi'_c$ as the maximum value of the radial component of the Lagrangian displacement at the flux tube boundary, which is linked to $\Xi_c$ by $\Xi'_c = \Xi_c J_1(n_R)$. We know that $\xi_c$ and $P'$ must be finite everywhere and non-zero at some point inside the flux tube, otherwise there would be no wave energy inside the flux tube. This happened in the previous limiting case (see Sect. 3.1), where in the long wavelength limit the external wave perturbations were exactly zero and therefore we found no external energy. A possible solution is

$$A_i = - \frac{\rho_{0,i} \left(\omega_{A,e}^2 - \omega_{A,e}^2\right)}{j_{0,1} J_1(j_{0,1})} R \Xi'_c,$$

$$A_e = - \frac{\rho_{0,e} \left(\omega_{A,e}^2 - \omega_{A,e}^2\right)}{(k_c R) J_1(k_c R)} R \Xi'_c = 0.$$

After some straightforward calculations we find that the energy inside the flux tube for fast sausage modes at the cut-off wavenumber in the zero plasma beta approximation is given by

$$\langle KE \rangle_i = \frac{D_{0j}}{4} \omega_{A,e}^2 R^2 \Xi_c^2,$$

$$\langle ME \rangle_i = \frac{D_{0j}}{4} \omega_{A,e}^2 R^2 \Xi_c^2,$$

$$\langle TE \rangle_i = 0,$$

$$\langle S \rangle_i = 2 \frac{D_{0j}}{4} \omega_{A,e}^2 c_{s1} R^2 \Xi_c^2 1_z,$$

$$\langle T \rangle_i = 0.$$  \hspace{1cm} (28)

Because we are working in the cold plasma approximation we expect no thermal energy and this is confirmed here. We note that there is no equipartition of energy between magnetic and kinetic energy. This is not a problem since we have not included the external energy at the moment. We have not mentioned the group speed since we have not yet calculated the external energy.
would expect a finite, but possibly large, external energy. We indeed found that for all trapped wave modes the energy is finite, but for $kr \rightarrow k_f R$ the external energy becomes much larger than the energy inside the flux tube. This confirms the mathematical result that at the cut-off wavenumber the external kinetic energy becomes infinite.

An infinitely large kinetic energy is not physically acceptable. The reason that the energy is infinite is quite simple and substituting $\omega = \omega_{\text{re}}(k_R)$ for all trapped wave modes the energy is finite, but for all trapped wave modes the energy is finite, we would expect a finite, but possibly large, external energy. We in-

The energy inside and outside the flux tube. For the energy inside the flux tube we have derived formulas in Sect. 3.2.1. To calculate the energy outside the flux tube we need to realize that the wave energy has only propagated to a distance $d = V \Delta t$, because the energy propagates at a velocity $V$. Therefore the integral we need to compute is not Eq. 18, but we should compute

$$\langle KE \rangle_c = 2 \pi \int_R^d r dr$$

(29)

since the energy is zero when going further from the flux tube than a distance $d$. This formula would give us a finite kinetic energy as we would expect. The shortcoming in the modelling of the wave modes is assuming that these waves exist over the entire finite domain, while in reality they only exist up to a certain distance $d$ outside the flux tube. We can summarise this particular wave mode as a case where no normal mode solution is found that is physically acceptable, but there is actually a wave mode solution to the initial value problem (Roberts & Boardman 1962).

The physical distance $d$ is almost impossible to quantitatively determine since the transport of energy is difficult to identify, even using the highest-resolution solar observations. However, there are ways to infer this distance $d$. We turn to the paper by Van Doorsselaere et al. (2014) and use the influence radius of a flux tube strand $d_i = R/\sqrt{f}$ to quantify the distance $d$. Using this influence radius $d_i$ in combination with Eq. 29 should result in finite values for the energy of the wave mode. The general calculation, when not using a specific wave mode and when integrating to a distance $d_i$, results in equations similar to Eqs. 19 with some adaptations, e.g.

$$\langle KE \rangle_c = C \int_R^\infty H_1^0(k_R) r dr,$$

where $C$ is a non-zero constant. The Hankel function squared can be written as $H_1^0(k_R) = J_1^2(k_R) + Y_1^2(k_R)$, where $J_1(.)$ and $Y_1(.)$ are Bessel functions of the first and second kind of order zero. This function is not square-integrable, i.e. the function is nonzero at infinity and therefore the integral results in an infinite kinetic energy.

Naturally, there is also a physical reason why the kinetic energy is infinite. Assume that at a certain time $t$ a flux tube in equilibrium is suddenly perturbed such that the excited wave mode is a fast sausage wave at the cut-off wavenumber. Because the wave mode is leaky, the energy will be transported outward from the flux tube at a constant speed. When we observe this wave mode at time $t + \Delta t$ and we want to calculate the energy, we need both the energy inside and outside the flux tube. For the energy inside the flux tube we have derived formulas in Sect. 3.2.1. To calculate the external energy we need to realize that the wave energy has only propagated to a distance $d = V \Delta t$, because the energy propagates at a velocity $V$. Therefore the integral we need to compute is not Eq. 18, but we should compute

$$\langle KE \rangle_c = 2 \pi \int_R^d \langle KE \rangle_c r dr$$

(30)

where $C_r$ and $C_z$ are given by Eq. 16 taken with values outside the flux tube. We have similar equations for the magnetic energy, the thermal energy, the total energy, the Poynting flux, the flux of thermal energy, and the flux of energy.

We can now proceed as in the case of the energy inside the flux tube for fast waves. We already know the complete solution to the dispersion relation and the amplitude $A_0$ (see Sect. 3.2.1). Therefore we can expand the Bessel functions for small arguments and using Eq. 30 we find that

$$\langle KE \rangle_c = \frac{\rho_0 c e_\omega^2}{4} \omega^{\text{re}}R^2 \Xi^2 \ln(1/f),$$

$$\langle ME \rangle_c = \frac{\rho_0 c e_\omega^2}{4} \omega^{\text{re}}R^2 \Xi^2 \ln(1/f),$$

$$\langle IE \rangle_c = 0,$$

$$\langle \mathcal{F} \rangle_c = 2 \frac{\rho_0 c e_\omega^2}{4} \omega^{\text{re}}R^2 \Xi^2 \ln(1/f) 1_z,$$

$$\langle \mathcal{F} \rangle_c = 0.$$
In Sect. 2 we have calculated the energy in wave modes using the complete equilibrium model, i.e. no additional assumptions were made. The resulting equations (Eqs. 15, 17, and 19) can be applied to both surface and body modes and the energy both inside and outside the flux tube can be calculated. As input for the calculations one needs the sound and Alfvén speeds, the phase speed of the wave, the plasma density, the radius of the flux tube, and the amplitude \( A \) that occurs in the mathematical description of the wave. The amplitude \( A \) can be linked with either the radial or longitudinal component of the Lagrangian displacement at the flux tube boundary (see Eqs. 16). All these parameters are available from observations and/or empirical models of the solar atmosphere. The phase speed of the wave is the hardest parameter to determine for observations in the lower solar atmosphere, since observations are only taken at one or a few heights in the magnetic structure. A method for determining the phase speed of waves using only intensity images at one height in the solar atmosphere has been described in Moreels et al. (2015). The key to determining the phase speed is to be able to identify the fractional changes in intensity and area, for which a high resolution is needed. Thus, as we move into the era of the Daniel K. Inouye Solar Telescope (DKIST, formerly known as the Advanced Technology Solar Telescope (ATST), Elmore et al. 2014) and the European Solar Telescope (EST, Collados et al. 2010), this paper combined with Moreels et al. (2015) will provide a way to analyse sausage modes in the lower solar atmosphere as the methods of observation become more advanced. It could also be said that the results of this work will become more accurate over time as the resolution of the data increases, therefore this work will become even more useful in the future.

In Sect. 3 and in Appendix A we have applied the general energy formulas to the case of sausage modes in thin flux tubes (in coronal/chromospheric flux tubes). We have formulas to calculate the energy for slow sausage modes in the long wavelength limit for both general plasma beta (Eqs. 24) and low plasma beta values (Eqs. 27). The slow sausage wave is dominated by the longitudinal component of the Lagrangian displacement and therefore the energy is expressed in terms of this displacement. In the long wavelength limit we find the equipartition between kinetic and potential energy. In the cold plasma approximation (i.e. the plasma beta is zero) the potential energy consists only of thermal energy and the magnetic energy is exactly zero. The group speed is equal to the internal tube speed, as expected. In Appendix A we list the energy equations for both the fundamental slow sausage body mode under coronal conditions and the sausage surface mode under photospheric conditions in thin flux tubes (i.e. \( kR \) is small but non-zero). In this section we also found the dispersive first-order correction to the phase speed in the thin tube limit for both the fundamental slow sausage body mode under coronal conditions and the sausage surface mode under photospheric conditions. From an observational point of view the most important quantity is the averaged flux of energy through a surface perpendicular to the flux tube \( \langle F_S \rangle \), which is given by

\[
\langle F_S \rangle = \frac{\rho_0}{2} V^2 c_{T,1} I_z. \tag{34}
\]

The flux \( \langle F_S \rangle \) is expressed in W/m² and can easily be calculated. In the above formula \( f \) is the filling factor which is available in most observations. The equilibrium density inside the flux tube \( \rho_0 \) can be found from empirical models of the solar atmosphere. We also have \( V' \), which is the longitudinal component of the velocity perturbation at the flux tube boundary, for on-disk obser-
vations this is mainly the line-of-sight velocity. Finally, $c_{T1}$ is the phase speed of the wave.

We applied the general energy equations to fast sausage modes at the cut-off wavenumber with the assumption that the plasma beta is zero. Here the wave mainly has a radial component of the Lagrangian displacement resulting in energy that depends on the radial component of the Lagrangian displacement. The energy inside the flux tube was easily calculated, but the energy outside the flux tube was more involved. We finally combined the concept of flux tube strand influence radii as explained in Van Doorsselaere et al. (2014) to arrive at energy formulas for the fast sausage wave (Eqs. 32). The energy equations also depend on the filling factor $f$ which is available from modern observations. We compared the energy in fast sausage waves with the energy in transverse kink waves and found that, depending on the amplitude of the radial component of the Lagrangian displacement, either sausage waves or transverse kink waves can contain more energy. As before we discuss the averaged flux of energy $\langle F_S \rangle$ in greater detail. We know that

$$\langle F_S \rangle = f(1 + \ln(1/f)) \frac{P_{0e}}{\omega_{\lambda e}^2} \xi c_{\lambda e} I_z.$$  \hspace{1cm} (35)

Again, the flux $\langle F_S \rangle$ is expressed in W/m² and can easily be calculated. The filling factor $f$ is available from observations. The equilibrium density outside the flux tube $P_{0e}$ can be found in empirical models of the solar atmosphere. We use $\xi$ to denote the radial component of the Lagrangian displacement at the flux tube boundary, which can be quantified using the radial change of the structure under observation. The speed $c_{\lambda e}$ is the phase speed of the wave. The frequency $\omega_{\lambda e}$ is the frequency of the wave.

Finally, we stress once more the applicability of these energy equations. The general energy equations are applicable for any axisymmetric magnetoacoustic wave in a cylindrical plasma structure. In this way the energy equations can be used in all regions of the solar atmosphere, from the photosphere up to the corona.

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References


Appendix A: Energy in slow wave modes in thin flux tubes

We now focus on the energy expressions for slow waves in thin flux tubes, i.e. with a small but non-zero value of $kR$. As in Sect. 3.1 we find that as $kR \to 0$ the frequency of slow modes is given by $\omega = \omega_{T,i} \pm (kR)^2 \ln(kR)\omega_1$, where the ‘plus’ is the fundamental body mode and the ‘minus’ is the surface mode. We have the same expressions for $n_iR$, $\kappa_iR$, and $\omega_1$ as in Sect. 3.1 and so we do not list them here. To calculate the amplitudes $A_i$ and $A_e$ we take some extra terms from the Bessel function expansions into account.

We know that

$$A_i = (kR) \ln(kR) \left( -2i \rho_{0,e} \omega_{T,i} \omega_1 (1 + c_{A,i}^2/c_{T,i}^2) \right) \left( 1 + \frac{(n_iR)^2}{2} \right) R \Xi'',$$

$$A_e = \frac{\rho_{0,e}}{\rho_{0,i}} \left( \frac{\omega_{T,i}^2 - \omega_{T,e}^2}{\omega_{T,i}^2 - \omega_{A,i}^2} \right) \kappa_i R K_i(k_iR)$$

$$= \frac{\rho_{0,e}}{2} \frac{\omega_{T,i}^2 - \omega_{T,e}^2}{\omega_{A,i}^2} R \Xi' (kR) \left( 1 - \frac{(n_iR)^2}{8} + \frac{(k_iR)^2}{8} \right),$$

where there are some extra terms from the Bessel function expansions. In the first line the ‘plus’ indicates surface modes, while the ‘minus’ indicates the fundamental body mode. To calculate the energy inside the flux tube (Eqs. 15 and 17) we also need to approximate some products of Bessel functions, i.e.

$$J_0(n_iR)^2 + J_1(n_iR)^2 = 1 \pm \frac{(n_iR)^2}{4}$$

$$J_1(n_iR)^2 - J_0(n_iR)J_2(n_iR) = \frac{(n_iR)^2}{8}$$

$$K_i(k_iR)^2 - K_0(k_iR)^2 = \frac{1}{(k_iR)^2}$$

$$K_0(k_iR)K_2(k_iR) - K_1(k_iR)^2 = -\ln((k_iR)^2/4) \approx -\ln(kR).$$

In the first line the ‘minus’ indicates the fundamental body mode, while the ‘plus’ indicates surface modes. These expansion can all be checked in Abramowitz & Stegun (1972). We now have all the ingredients needed to calculate the energy in slow sausage wave modes in thin flux tubes. The energy inside the flux tube is given by

$$\langle KE \rangle_i = \frac{\rho_{0,i}}{4} \omega_{T,i}^2 \pi R^2 \Xi^2 \left( 1 \pm 3 \frac{\rho_{0,i}}{\rho_{0,e}} \left( \frac{\omega_{A,i}^2 - \omega_{T,i}^2}{\omega_{T,i}^2 - \omega_{A,i}^2} \right) \ln(kR) + \frac{\omega_{T,i}^4}{\omega_{A,i}^4} (kR)^2 \right),$$

$$\langle ME \rangle_i = \frac{\rho_{0,i}}{4} \omega_{T,i}^2 \pi R^2 \Xi^2 \left( 1 \pm 3 \frac{\rho_{0,i}}{\rho_{0,e}} \left( \frac{\omega_{A,i}^2 - \omega_{T,i}^2}{\omega_{T,i}^2 - \omega_{A,i}^2} \right) \ln(kR) + \frac{\omega_{T,i}^2}{\omega_{A,i}^2} (kR)^2 \right),$$

$$\langle IE \rangle_i = \frac{\rho_{0,i}}{4} \omega_{T,i}^2 \pi R^2 \Xi^2 \left( 1 \pm 3 \frac{\rho_{0,i}}{\rho_{0,e}} \left( \frac{\omega_{A,i}^2 - \omega_{T,i}^2}{\omega_{T,i}^2 - \omega_{A,i}^2} \right) \ln(kR) \right),$$

$$\langle TE \rangle_i = \frac{\rho_{0,i}}{2} \omega_{T,i}^2 \pi R^2 \Xi^2 \left( 1 \pm 3 \frac{\rho_{0,i}}{\rho_{0,e}} \left( \frac{\omega_{A,i}^2 - \omega_{T,i}^2}{\omega_{T,i}^2 - \omega_{A,i}^2} \right) \ln(kR) \right),$$

$$\langle U \rangle_i = \frac{\rho_{0,i}}{2} \omega_{T,i}^2 \pi R^2 \Xi^2 c_T \frac{\omega_{T,i}^2}{\omega_{A,i}^2} (kR)^2 1_z,$$

$$\langle T \rangle_i = \frac{\rho_{0,i}}{2} \omega_{T,i}^2 \pi R^2 \Xi^2 c_T \left( 1 \pm 3 \frac{\rho_{0,i}}{\rho_{0,e}} \left( \frac{\omega_{A,i}^2 - \omega_{T,i}^2}{\omega_{T,i}^2 - \omega_{A,i}^2} \right) \ln(kR) \right) 1_z,$$

$$\langle E \rangle_i = \frac{\rho_{0,i}}{2} \omega_{T,i}^2 \pi R^2 \Xi^2 c_T \left( 1 \pm 3 \frac{\rho_{0,i}}{\rho_{0,e}} \left( \frac{\omega_{A,i}^2 - \omega_{T,i}^2}{\omega_{T,i}^2 - \omega_{A,i}^2} \right) \ln(kR) + \frac{\omega_{T,i}^2}{\omega_{A,i}^2} (kR)^2 \right) 1_z,$$

(A.1)

where we have kept all terms of order $(kR)^2$ or lower. The ‘minus’ indicates the fundamental body mode, while the ‘plus’ indicates surface modes. We notice the equipartition between kinetic and potential (i.e. the sum of both magnetic and thermal energy) energy.
The energy outside the flux tube (Eqs. 19) can also be calculated and we find

\[
\langle KE \rangle_e = \frac{\rho_0 e^2}{4} \omega_T^2 \pi R^2 \xi_e \left( \frac{\omega_T^2}{\omega_{\lambda,i}^2} \right) \left( kR \right)^2 \left( -\ln(kR) \right) + \left( \frac{1}{2} \frac{\rho_0}{\rho_0 e^2} \frac{\omega_{T,1}^2}{\omega_{T,1}^2} \frac{\omega_{\lambda,i}^2}{\omega_{\lambda,i}^2} \frac{\omega_{T,e}^2 + \omega_{\lambda,e}^2}{\omega_{T,e}^2 + \omega_{\lambda,e}^2} \frac{\omega_{T,1}^2 - \omega_{T,1}^2}{\omega_{T,1}^2 - \omega_{T,1}^2} \right) (kR)^2.
\]

\[
\langle ME \rangle_e = \frac{\rho_0 e^2}{4} \omega_T^2 \pi R^2 \xi_e \left( \frac{\omega_T^2}{\omega_{\lambda,i}^2} \right) \left( kR \right)^2 \left( -\ln(kR) \right) + \left( \frac{1}{2} \frac{\rho_0}{\rho_0 e^2} \frac{\omega_{T,1}^2}{\omega_{T,1}^2} \frac{\omega_{\lambda,i}^2}{\omega_{\lambda,i}^2} \frac{\omega_{T,e}^2 + \omega_{\lambda,e}^2}{\omega_{T,e}^2 + \omega_{\lambda,e}^2} \frac{\omega_{T,1}^2 - \omega_{T,1}^2}{\omega_{T,1}^2 - \omega_{T,1}^2} \right) (kR)^2.
\]

\[
\langle IE \rangle_e = \frac{\rho_0 e^2}{4} \omega_T^2 \pi R^2 \xi_e \left( \frac{\omega_T^2}{\omega_{\lambda,i}^2} \right) \left( kR \right)^2 \left( -\ln(kR) \right) + \left( \frac{1}{2} \frac{\rho_0}{\rho_0 e^2} \frac{\omega_{T,1}^2}{\omega_{T,1}^2} \frac{\omega_{\lambda,i}^2}{\omega_{\lambda,i}^2} \frac{\omega_{T,e}^2 + \omega_{\lambda,e}^2}{\omega_{T,e}^2 + \omega_{\lambda,e}^2} \frac{\omega_{T,1}^2 - \omega_{T,1}^2}{\omega_{T,1}^2 - \omega_{T,1}^2} \right) (kR)^2,
\]

\[
\langle TE \rangle_e = \langle KE \rangle_e + \langle ME \rangle_e + \langle IE \rangle_e,
\]

\[
\langle S \rangle_e = \frac{\rho_0 e^2}{2} \omega_T^2 \pi R^2 \xi_e \left( \frac{\omega_T^2}{\omega_{\lambda,i}^2} \right) \left( kR \right)^2 \left( -\ln(kR) \right) + \left( \frac{1}{2} \frac{\rho_0}{\rho_0 e^2} \frac{\omega_{T,1}^2}{\omega_{T,1}^2} \frac{\omega_{\lambda,i}^2}{\omega_{\lambda,i}^2} \frac{\omega_{T,e}^2 + \omega_{\lambda,e}^2}{\omega_{T,e}^2 + \omega_{\lambda,e}^2} \frac{\omega_{T,1}^2 - \omega_{T,1}^2}{\omega_{T,1}^2 - \omega_{T,1}^2} \right) (kR)^2.
\]

\[
\langle T \rangle_e = \frac{\rho_0 e^2}{2} \omega_T^2 \pi R^2 \xi_e \left( \frac{\omega_T^2}{\omega_{\lambda,i}^2} \right) \left( kR \right)^2 \left( -\ln(kR) \right) + \left( \frac{1}{2} \frac{\rho_0}{\rho_0 e^2} \frac{\omega_{T,1}^2}{\omega_{T,1}^2} \frac{\omega_{\lambda,i}^2}{\omega_{\lambda,i}^2} \frac{\omega_{T,e}^2 + \omega_{\lambda,e}^2}{\omega_{T,e}^2 + \omega_{\lambda,e}^2} \frac{\omega_{T,1}^2 - \omega_{T,1}^2}{\omega_{T,1}^2 - \omega_{T,1}^2} \right) (kR)^2.
\]

\[
\langle E \rangle_e = \langle S \rangle_e + \langle T \rangle_e,
\]

where again we have kept all terms of order \((kR)^2\) or lower. The dominant term in these energy equations is the \((kR)^2 \left( -\ln(kR) \right)\) term. This shows that we do not find equipartition in this case since the coefficients in the kinetic and magnetic energy before the \((kR)^2 \left( -\ln(kR) \right)\) term are not the same. This puzzling result was further investigated numerically. We used different equilibrium parameters with different small values of \(kR\). We discovered a small error in the equipartition of energy of the order of \(10^{-2}\) of the total energy. When using the full set of equations to calculate the energy (i.e. Eqs. 17 and 19) we did find equipartition between kinetic and potential energy. This clearly shows that there is indeed equipartition, but approximating the Bessel functions has introduced a small error.

When taking the long wavelength limit (i.e. \(kR = 0\)) we find that Eqs. A.1 and A.2 simplify to Eqs. 24 as expected.