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WASP-10b: a 3M$_J$, gas-giant planet transiting a late-type K star

D. J. Christian,¹,²⋆ N. P. Gibson,¹ E. K. Simpson,¹ R. A. Street,¹,³ I. Skillen,⁴
D. Pollacco,¹ A. Collier Cameron,⁵ Y. C. Joshi,¹ F. P. Keenan,¹ H. C. Stempels,⁵
C. A. Haswell,⁶ K. Horne,⁵ D. R. Anderson,⁷ S. Bentley,⁷ F. Bouchy,⁸,⁹
W. I. Clarkson,⁶,¹⁰ B. Enoch,⁶ L. Hebb,⁵ G. Hébrard,⁸ C. Hellier,⁷ J. Irwin,¹¹
S. R. Kane,¹² T. A. Lister,³,⁵,⁷ B. Loeillet,¹³ P. Maxted,⁷ M. Mayor,¹⁴ I. McDonald,⁷
C. Moutou,¹³ A. J. Norton,⁶ N. Parley,⁶ F. Pont,¹⁴,¹⁵ D. Queloz,¹⁴ R. Ryans,¹
B. Smalley,⁷ A. M. S. Smith,⁵ I. Todd,¹ S. Udry,¹⁴ R. G. West,¹⁶ P. J. Wheatley¹⁷
and D. M. Wilson⁷

¹Astrophysics Research Centre, School of Mathematics and Physics, Queen’s University, University Road, Belfast, BT7 1NN
²Department of Physics and Astronomy, California State University Northridge, 18111 Nordhoff Street, Northridge, CA 91330-8268, USA
³Las Cumbres Observatory, 6740 Cortona Dr Suite 102, Santa Barbara, CA 93117, USA
⁴Isaac Newton Group of Telescopes, Apartado de Correos 321, E-38790 Santa Cruz, de la Palma, Tenerife, Spain
⁵School of Physics and Astronomy, University of St Andrews, North Haugh, St Andrews, Fife KY16 9SS
⁶Department of Physics and Astronomy, The Open University, Milton Keynes, MK7 6AA
⁷Astrophysics Group, Keele University, Staffordshire, ST5 5BG
⁸Institut d’Astrophysique de Paris, CNRS (UMR 7095) – Université Pierre and Marie Curie, 98bis bvd. Arago, 75014 Paris, France
⁹Observatoire de Haute-Provence, 04870 St Michel l’Observatoire, France
¹⁰STScI, 3700 San Martin Drive, Baltimore, MD 21218, USA
¹¹Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA 02138, USA
¹²Michelson Science Center, Caltech, MS 100-22, 770 South Wilson Avenue Pasadena, CA 91125, USA
¹³Laboratoire d’Astrophysique de Marseille, BP 8, 13376 Marseille Cedex 13, France
¹⁴Observatoire de Genève, Université de Genève, 51 Ch. des Maillettes, 1290 Sauverny, Switzerland
¹⁵School of Physics, University of Exeter, Stocker Road, Exeter, EX4 4QL
¹⁶Department of Physics and Astronomy, University of Leicester, Leicester, LE1 7RH
¹⁷Department of Physics, University of Warwick, Coventry, CV4 7AL

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ABSTRACT

We report the discovery of WASP-10b, a new transiting extrasolar planet (ESP) discovered by the Wide Angle Search for Planets (WASP) Consortium and confirmed using Nordic Optical Telescope Fibre-fed Echelle Spectrograph and SOPHIE radial velocity data. A 3.09-d period, 29 mmag transit depth and 2.36 h duration are derived for WASP-10b using WASP and high-precision photometric observations. Simultaneous fitting to the photometric and radial velocity data using a Markov Chain Monte Carlo procedure leads to a planet radius of 1.28 R$_J$, a mass of 2.96 M$_J$ and eccentricity of ≈0.06. WASP-10b is one of the more massive transiting ESPs, and we compare its characteristics to the current sample of transiting ESP, where there is currently little information for masses greater than ≈2 M$_J$ and non-zero eccentricities. WASP-10’s host star, GSC 2752−00114 (USNO-B1.0 1214−0586164) is among the fainter stars in the WASP sample, with V = 12.7 and a spectral type of K5. This result shows promise for future late-type dwarf star surveys.

Key words: methods: data analysis – techniques: photometric – techniques: radial velocities – stars: planetary systems.

1 INTRODUCTION

Photometric transit observations of extrasolar planets (ESP) are important because the transit strongly constrains their orbital inclination and allows accurate physical parameters for the planet to be
derived. Their mass–radius relation allows us to probe their internal structure and is vital to our understanding of orbital migration and planetary formation. The radial velocity (RV) measurements which are used to confirm a candidate transiting ESP also provide more complete information on the orbital eccentricity.

As wide-field photometric transit surveys have collected additional sky and temporal coverage, and understood their noise components (Collier Cameron et al. 2007a; Smith et al. 2007), the number of transiting ESP has grown to over 50 in line with earlier predictions (Horne 2003). Recently, one such survey SuperWASP (Pollacco et al. 2006) published its first five confirmed ESP, all of which have periods of less than 3 d (Anderson et al. 2008; Collier Cameron et al. 2007b; Pollacco et al. 2008; Wilson et al. 2008), and reported an additional 10\(^2\) (Helling et al. 2008; Hebb et al. 2008; Joshi et al. 2008; West et al. 2008). SuperWASP is performing a 'shallow-but-wide' transit search, designed to find planets that are not only sufficiently bright (\(9 < V < 13\)) for high-precision RV follow-up to be feasible on telescopes of moderate aperture, but also for detailed studies such as transmission spectroscopy during transits. Details of the WASP project and observatory infrastructure are described in Pollacco et al. (2006).

In this paper, we present the WASP photometry of 1SWASP J231558.30+312746.4 (GSC 2752−00114), higher precision photometric follow-up observations with the MERCATOR and Tenagra telescopes and high-precision RV observations with the Nordic Optical Telescope new Fibre-fed Echelle Spectrograph (FIES) and the Observatoire de Haute-Provence (OHP) SOPHIE collaboration. These observations lead to the discovery and confirmation of a new, relatively high-mass, gas-giant exoplanet, WASP-10b.

2 OBSERVATIONS

1SWASP J231558.30+312746.4 (GSC 2752−00114) was monitored by SuperWASP-N starting in 2004. SuperWASP is a multi-camera telescope system with SuperWASP-North located in La Palma and consisting of 8 Canon 200-mm f/1.8 lenses each coupled to e2v 2048 × 2048 pixel back illuminated CCDs. This combination of lens and camera yields a field of view of 7.8 x 7.8 with an angular size of 14.2 arcsec per pixel. During 2004, SuperWASP was run with four or five cameras as the operations moved from commissioning to routine automated observing. We show the WASP-10b SuperWASP light curve in Fig. 1.

2.1 Higher precision photometry

We obtained photometry of WASP-10 with the MEROPE instrument on the 1.2 m MERCATOR Telescope in V band on 2007 September 1. Only a partial transit was observed due to uncertainties in the period and epoch from the SuperWASP data. Observations were in the V band with 2 x 2 binning over ~2.9 h. Despite clear conditions, exposure times were varied from 25 to 30 s to account for changes in seeing and to keep below the saturation limits of the chip. This allowed 170 images to be taken. There was a drift of only 1 binned pixel in x and y on the chip during the run. The MEROPE images were first de-biased and flat-fielded with combined twilight flats using IRAF and the APHOT package to obtain aperture photometry of the target and five nearby companion stars using a 10 pixel radius. Finally, the light curve was extracted and normalized to reveal a depth of ~33 mmag.

Further observations of WASP-10 were taken as part of an observing program sponsored by the Las Cumbres Observatory Global Telescope Network\(^2\) on the Tenagra II, 0.81 m F7 Ritchey–Chretien telescope sited in the Sonora desert in S. Arizona, USA. The science camera contains a 1 x 1 k SITe CCD with a pixel scale of 0.87 arcsec pixel\(^{-1}\) and a field of view of 14.8 x 14.8 arcmin. The filter set is the standard Johnson/Cousins/Bessel UBVR set and the data presented here have been taken in V band.

Calibration frames were obtained automatically every twilight, and the data were de-biased and flat-fielded using the calibration section of the SuperWASP pipeline. Object detection and aperture photometry were then performed using DAOPHOT (Stetson 1987) within IRAF. Differential photometry was derived from a selection of typically 5 to 10 comparison stars within the frame.

These confirmed that the object had a sharp egress with an amplitude of 0.033 ± 0.001 mag. The MERCATOR V and Tenagra I light curves show consistent transit depths, confirming that the companion is a transiting ESP.

2.2 Spectroscopic follow-up

We obtained high-precision RV follow-up observations of WASP-10 with the 2.5 m Nordic Optical Telescope (NOT) new FIES, supplemented with observations from the Observatoire de Haute-Provence’s 1.93 m telescope and the SOPHIE spectrograph (Bouchy et al. 2006). We present a summary of the FIES and SOPHIE RV data in Table 1.

2.2.1 NOT and FIES

Spectroscopic observations were obtained using the new FIES spectrograph mounted on the NOT Telescope. A total of seven RV points were obtained during 2007 December 2, 28–31 and 2008 January 24–25. WASP-10 required observations with an exposure time of 2400 s due its relative faintness (V = 12.7) yielding a peak signal-to-noise ratio per resolution element of ~60–70 in the H\alpha region. FIES was used in medium resolution mode with R = 46 000 with simultaneous ThAr calibration. We used the bespoke data reduction package FIESTOOL\(^3\) to extract the spectra and a specially developed IDL line-fitting code to obtain RVs with a precision of 15–25 m s\(^{-1}\) (except for the poor night of 2008 January 24, JD 2454490).

2.2.2 OHP 1.9 m and SOPHIE

Additional RV measurements were taken for WASP-10 on 29 and 30 August 2007, and again between 2008 Feb 11 and 15 with

\(^1\) http://www.inscience.ch/transits/

\(^2\) http://www.lcogt.net

\(^3\) http://www.not.iac.es/instruments/fies/fiestool/FIEStool.html
the 2 to 3 m s$^{-1}$ photon noise. Thus, the average RV error is basis. Although errors for each RV measurement are limited by the

\[ \text{RV drift never exceeded 2 to 3 m s}^{-1}. \]

were carefully removed during the process. This merged spectrum was then continuum normalized with a low-order polynomial to retain the shape of the broadest spectral features. The total signal-to-noise ratio estimates for each spectra were 40 000. Thorium–Argon calibration images were taken at the start and end of each night, and at 2- to 3-h intervals throughout the night.

As previously undertaken for our analysis of WASP-1 (Stempels et al. 2007) and WASP-3 (Pollacco et al. 2008), we employed the methodology of Valenti & Fischer (2005), using the same tools, techniques and model atmosphere grid. We used the package Spectroscopy Made Easy (SME) (Valenti & Piskunov 1996), which combines spectral synthesis with multidimensional $\chi^2$ minimization to determine which atmospheric parameters best reproduce the observed spectrum of WASP-10 (effective temperature $T_{\text{eff}}$, surface gravity log $g$, metallicity [M/H], projected RV $v\sin i$, systemic RV $v_{\text{rad}}$, microturbulence $v_{\text{mic}}$ and the macroturbulence $v_{\text{mac}}$).

The four spectral regions we used in our analysis are: (i) 5160–5190 Å, covering the gravity-sensitive Mg b triplet; (ii) 5850–5950 Å, with the temperature and gravity-sensitive Na I D doublet; (iii) 6000–6210 Å, containing a wealth of different metal lines, providing leverage on the metallicity, and (iv) 6520–6600 Å, covering the strongly temperature-sensitive Hα line. In addition, we analysed a small region around the Li I 6708 line to possibly derive a lithium abundance, but no Li I 6708 was detected for WASP-10.

We find WASP-10b to have a radius $1.28^{+0.03}_{-0.03}$ km s$^{-1}$. We were not able to include the SOPHIE spectra in this analysis, which is known to suffer from problems with removal of the blaze function.

### 3 RESULTS AND ANALYSIS

#### 3.1 Stellar parameters

We merged all available WASP-10 FIES spectra into one high-quality spectrum in order to perform a detailed spectroscopic analysis of the stellar atmospheric properties. Radial velocity signatures were carefully removed during the process. This merged spectrum was then continuum normalized with a low-order polynomial to retain the shape of the broadest spectral features. The total signal-to-noise ratio of the combined spectrum was $\approx 180$ per pixel. We were not able to include the SOPHIE spectra in this analysis, because these spectra were obtained in the high-efficiency (HE) mode, which is known to suffer from problems with removal of the blaze function.

### Table 1. Journal of RV measurements for WASP-10 (1SWASP J231558.30+312746.4, USNO-B1.0 1214—0586164). Stellar coordinates are for the photometric apertures; the USNO-B1.0 number denotes the star for which the RV measurements were secured. The quoted uncertainties in the RV errors include components due to photon noise (Section 2.2) and 10 m s$^{-1}$ of jitter (Section 3.2) added in quadrature.

<table>
<thead>
<tr>
<th>BJD</th>
<th>$t_{\text{exp}}$</th>
<th>$V_{r}$ km s$^{-1}$</th>
</tr>
</thead>
<tbody>
<tr>
<td>FIES</td>
<td></td>
<td></td>
</tr>
<tr>
<td>245437.540</td>
<td>2400</td>
<td>$-11.028 \pm 0.026$</td>
</tr>
<tr>
<td>245446.377</td>
<td>2400</td>
<td>$-11.941 \pm 0.030$</td>
</tr>
<tr>
<td>245446.342</td>
<td>2400</td>
<td>$-11.003 \pm 0.018$</td>
</tr>
<tr>
<td>245446.335</td>
<td>2400</td>
<td>$-11.804 \pm 0.021$</td>
</tr>
<tr>
<td>245449.329</td>
<td>2400</td>
<td>$-11.013 \pm 0.143$</td>
</tr>
<tr>
<td>245449.358</td>
<td>2400</td>
<td>$-10.990 \pm 0.120$</td>
</tr>
<tr>
<td>245449.340</td>
<td>2400</td>
<td>$-11.955 \pm 0.024$</td>
</tr>
<tr>
<td>SOPHIE</td>
<td></td>
<td></td>
</tr>
<tr>
<td>245434.569</td>
<td>3300</td>
<td>$-11.657 \pm 0.008$</td>
</tr>
<tr>
<td>245434.505</td>
<td>3600</td>
<td>$-11.575 \pm 0.011$</td>
</tr>
<tr>
<td>245450.268</td>
<td>1680</td>
<td>$-11.027 \pm 0.014$</td>
</tr>
<tr>
<td>245450.268</td>
<td>1680</td>
<td>$-11.336 \pm 0.017$</td>
</tr>
<tr>
<td>245451.276</td>
<td>1680</td>
<td>$-11.990 \pm 0.020$</td>
</tr>
<tr>
<td>245451.262</td>
<td>1680</td>
<td>$-11.335 \pm 0.016$</td>
</tr>
<tr>
<td>245451.262</td>
<td>1680</td>
<td>$-11.244 \pm 0.016$</td>
</tr>
</tbody>
</table>

#### 3.2 Markov Chain Monte Carlo analysis

Transit timing and the RV measurements provide detailed information about the orbit. We modelled WASP-10b’s transit photometry and the reflex motion of the host star simultaneously using the Markov Chain Monte Carlo (MCMC) algorithm described in detail by Collier Cameron et al. (2007a), and the same techniques that were applied to WASP-3 by Pollacco et al. (2008) to which we refer the reader for more details.

We find WASP-10b to have a radius $1.28^{+0.03}_{-0.03}$ R$_J$, mass of $2.96^{+0.22}_{-0.17}$ M$_J$ and a significant non-zero eccentricity of $0.059^{+0.014}_{-0.004}$. The best-fitting solution for the MCMC model for a circular orbit ($e = 0$) has a $\chi^2$ 55 higher than the solution with a non-zero eccentricity, and thus, the eccentricity is significant at $> 99.6$ per cent confidence level using the F test. The values of the parameters of the optimal solution are given, together with their associated 1σ confidence intervals, in Table 3. The FIES + SOPHIE RV data measurements are plotted in Fig. 2 together with the best-fitting global fit to the SuperWASP-N, MERCATOR and Tenagra transit photometry.
3.3 Line-bisector variation

Line bisectors have been shown to be a powerful diagnostic in distinguishing true ESPs from blended and eclipsing stellar systems chromospheric activity (Queloz et al. 2001). Torres et al. (2004) showed that for OGLE-TR-33 line asymmetries which changed with a 1.95-d period, it was a blended system. From the cross-correlation function (CCF), we obtained the line bisectors and these are plotted, as a function of RV, in Fig. 3.

We determined the significance of the bisector variation as follows. We determined the inverse-variance-weighted averages of the RV and bisector span as

\[ \hat{v} = \frac{\sum v_i w_i}{\sum w_i}, \hat{b} = \frac{\sum b_i w_i}{\sum w_i}, \]

where the \( v_i \) and \( b_i \) are the RV and span bisector values, respectively, and the weights \( w_i \) are the inverse variances of the individual bisector measurements. The uncertainty in the span bisector is assumed to be 2.5 times the uncertainty on the RV in our data. If we define \( x_i = v_i - \hat{v} \) and \( y_i = b_i - \hat{b} \), then the slope is determined as

\[ \hat{a} = \frac{\sum x_i y_i w_i}{\sum x_i^2 w_i}; \text{Var}(\hat{a}) = \frac{1}{\sum x_i^2 w_i}. \]

The value of the scaling factor \( \hat{a} \) is determined with a signal-to-noise ratio

\[ \text{SNR} = \frac{\hat{a}}{\sqrt{\text{Var}(\hat{a})}}. \]

Table 3. WASP-10 system parameters and 1σ error limits derived from MCMC analysis.

<table>
<thead>
<tr>
<th>Parameter</th>
<th>Symbol</th>
<th>Value</th>
</tr>
</thead>
<tbody>
<tr>
<td>Transit epoch (BJD)</td>
<td>( T_0 )</td>
<td>( 2454357.8508_{-0.0003}^{+0.0004} ) d</td>
</tr>
<tr>
<td>Orbital period</td>
<td>( P )</td>
<td>( 3.0927636_{-0.000009}^{+0.000021} ) d</td>
</tr>
<tr>
<td>Planet/star area ratio</td>
<td>((R_p/R_*)^2)</td>
<td>(0.029_{-0.001}^{+0.001})</td>
</tr>
<tr>
<td>Transit duration</td>
<td>( \tau )</td>
<td>(0.09818_{-0.0019}^{+0.0019} ) d</td>
</tr>
<tr>
<td>Impact parameter</td>
<td>( \beta )</td>
<td>(0.568_{-0.084}^{+0.054} R_\odot )</td>
</tr>
<tr>
<td>Stellar reflex velocity</td>
<td>( K_1 )</td>
<td>(5.201_{-0.0084}^{+0.0084} ) km s(^{-1})</td>
</tr>
<tr>
<td>Centre-of-mass velocity</td>
<td>( \gamma )</td>
<td>(-11.485_{-0.0012}^{+0.0012} ) km s(^{-1})</td>
</tr>
<tr>
<td>Orbital semimajor axis</td>
<td>( a )</td>
<td>(0.0369_{-0.0004}^{+0.0004} ) AU</td>
</tr>
<tr>
<td>Orbital inclination</td>
<td>( i )</td>
<td>(86.9_{-0.5}^{+0.5} ) degrees</td>
</tr>
<tr>
<td>Orbital eccentricity</td>
<td>( \epsilon )</td>
<td>(0.059_{-0.004}^{+0.004})</td>
</tr>
<tr>
<td>Arg. periastron</td>
<td>( \omega )</td>
<td>(2.917_{-0.245}^{+0.222} ) (rad)</td>
</tr>
<tr>
<td>Stellar mass</td>
<td>( M_* )</td>
<td>(0.703_{-0.085}^{+0.068} M_\odot )</td>
</tr>
<tr>
<td>Stellar radius</td>
<td>( R_* )</td>
<td>(0.775_{-0.043}^{+0.043} R_\odot )</td>
</tr>
<tr>
<td>Stellar surface gravity</td>
<td>( \log g_* )</td>
<td>(4.51_{-0.05}^{+0.06} ) (CGS)</td>
</tr>
<tr>
<td>Stellar density</td>
<td>( \rho_* )</td>
<td>(1.51_{-0.20}^{+0.25} ) ( \rho_\odot )</td>
</tr>
<tr>
<td>Planet radius</td>
<td>( R_p )</td>
<td>(1.28_{-0.091}^{+0.077} R_J )</td>
</tr>
<tr>
<td>Planet mass</td>
<td>( M_p )</td>
<td>(2.96_{-0.17}^{+0.22} M_J )</td>
</tr>
<tr>
<td>Planetary surface gravity</td>
<td>( \log g_p )</td>
<td>(3.62 \pm 0.06 ) (CGS)</td>
</tr>
<tr>
<td>Planet density</td>
<td>( \rho_p )</td>
<td>(1.43_{-0.29}^{+0.31} ) ( \rho_J )</td>
</tr>
<tr>
<td>Planet temp ((A = 0))</td>
<td>( T_{\text{eq}} )</td>
<td>(1119_{-28}^{+28} ) K</td>
</tr>
<tr>
<td>Photometric data points</td>
<td>( \chi^2_{\text{phot}} )</td>
<td>4145</td>
</tr>
<tr>
<td>Spectroscopic data points</td>
<td>( \chi^2_{\text{spec}} )</td>
<td>17.2</td>
</tr>
<tr>
<td>Photometric data points</td>
<td>( N_{\text{phot}} )</td>
<td>4151</td>
</tr>
<tr>
<td>Spectroscopic data points</td>
<td>( N_{\text{spec}} )</td>
<td>14</td>
</tr>
</tbody>
</table>

We obtain SNR = 1.16 indicating a non-significant correlation between the bisector span and RV variations. This demonstrates that the CCF remains symmetric, and that the RV variations are not likely to be caused by line-of-site binarity or stellar activity and indicate WASP-10b is an exoplanet.
studied to rule out a putative outer plane that may be driving its eccentricity. Thus, the ≈6 per cent eccentricity of WASP-10b makes it an attractive target for future transit-timing variation studies, and for longer term RV monitoring to establish the mass and period of the putative outer planet.

The majority of transiting ESP found have masses below 1.5\(M_J\), although there are a few more massive ESP. HD 17156 and COROT-Exo-2 have similar masses to WASP-10b and although there are two more massive ESP, the nearly 9 \(M_J\) HAT P-2 (HD 147506b) (Bakos et al. 2007) and 7.3 \(M_J\) WASP-14b (Joshi et al. 2008), this higher mass region has been poorly explored. Additional transiting objects in the mass range are important for completing the current ESP mass–radius relations and constraining their compositions. The current sample of transiting ESP reveals a large range of densities. We derive a mean density for WASP-10b of ≈1.89 g cm\(^{-3}\) (1.42 \(\rho_J\)) and it would lie along the higher density contour in a mass–radius plot (Pollacco et al. 2008; Sozzetti et al. 2007).

One ultimate goal of our transit-search programme is to provide the observational grist that will stimulate and advance refined models for the formation and evolution of the hot and very hot Jupiters (e.g. Burrows et al. 1997; Fortney, Marley & Barnes 2007; Seager et al. 2007). By thus constraining the underlying physics, we will have a richer context for the interpretation of the lower mass planets expected from missions such as COROT and Kepler.

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